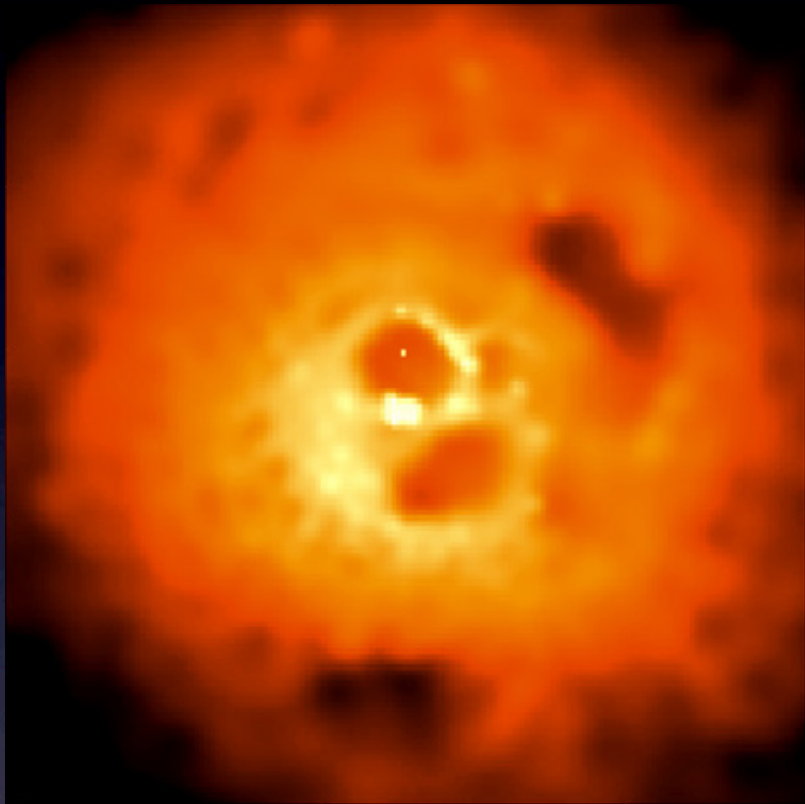


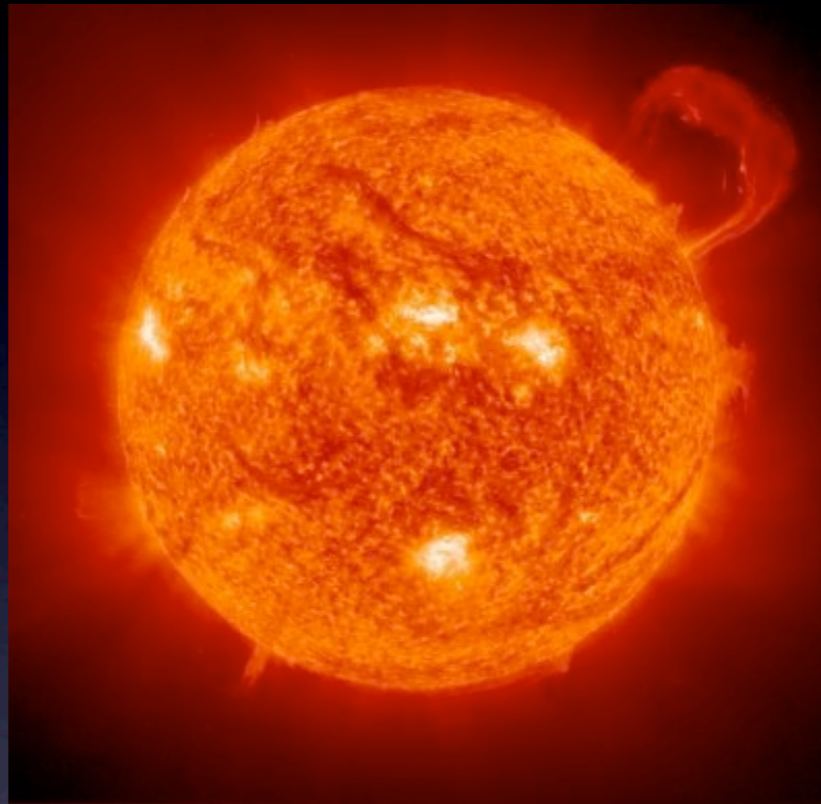
From Accretion Flow to Particle Acceleration: New Relativistic MHD jet solutions including Gravity

Peter Polko
Sera Markoff
David Meier

Introduction



Credit: NASA/IOA/A. Fabian et al.



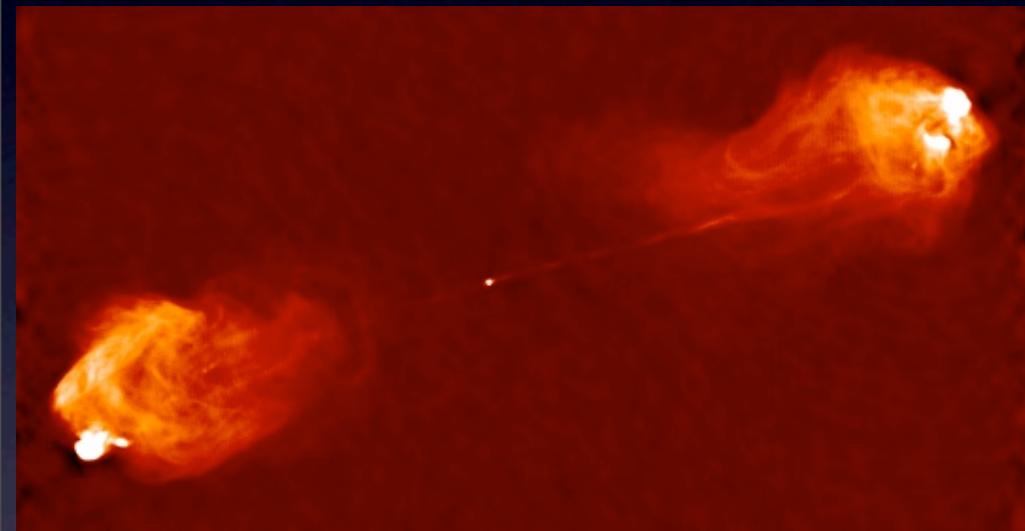
Credit: SOHO/EIT consortium

Perseus cluster

Young stellar object

Gamma-ray flashes

Jets!



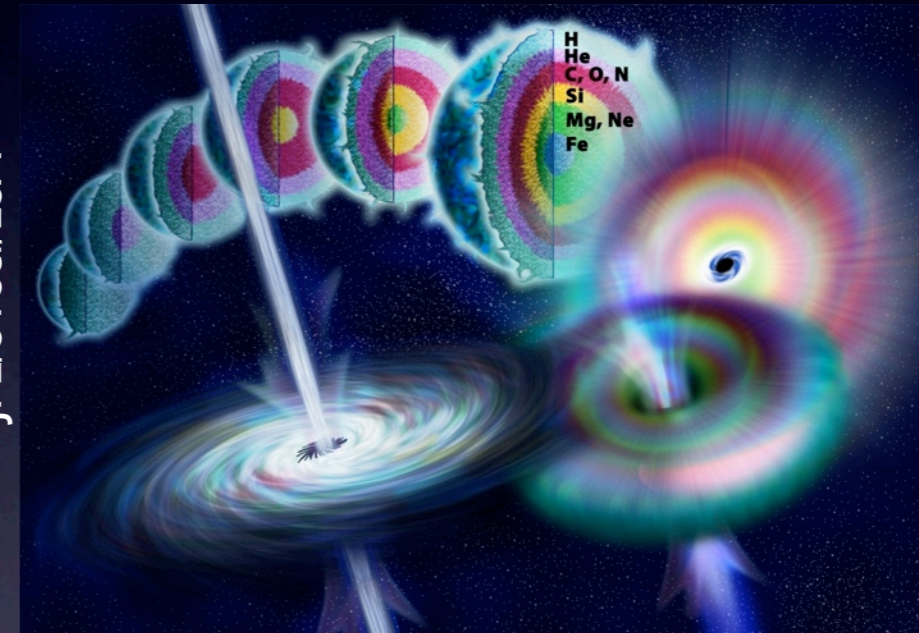
Credit: C. Carilli

Cygnus A



Credit: Universidad Nacional Autonoma de Maxico/
JPL/STSci/ESA

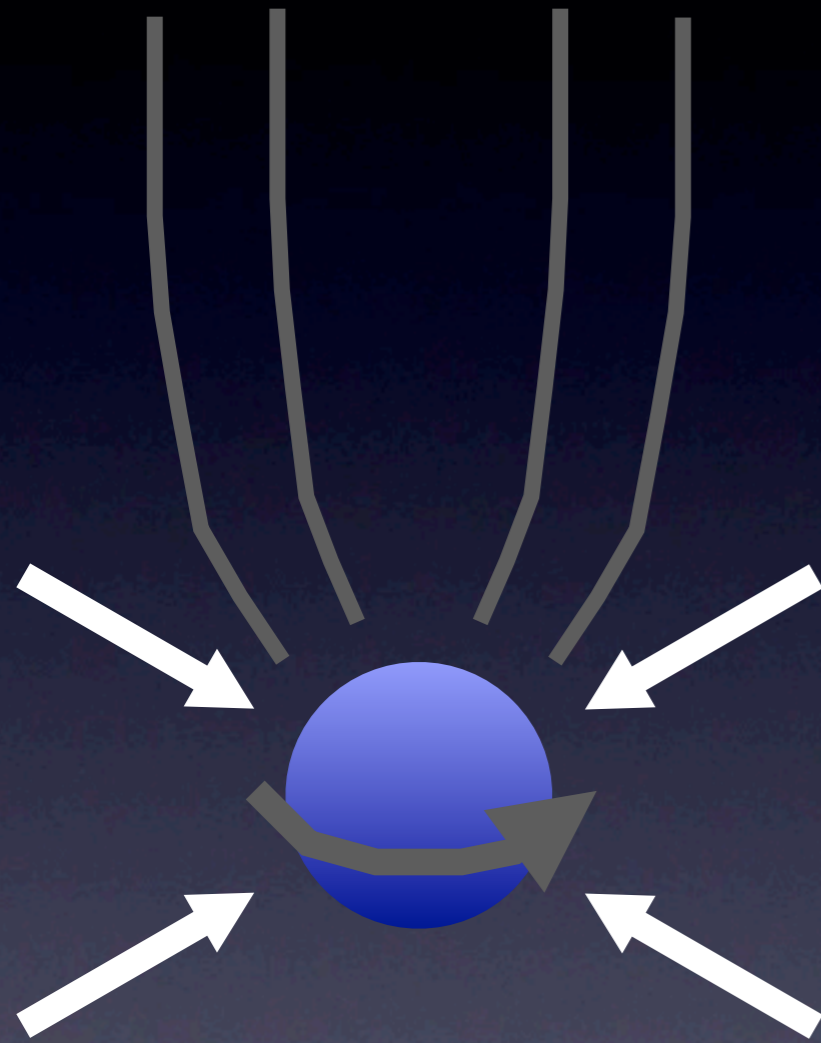
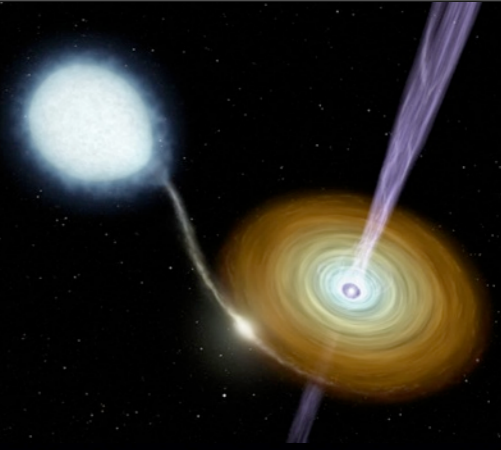
HH 30



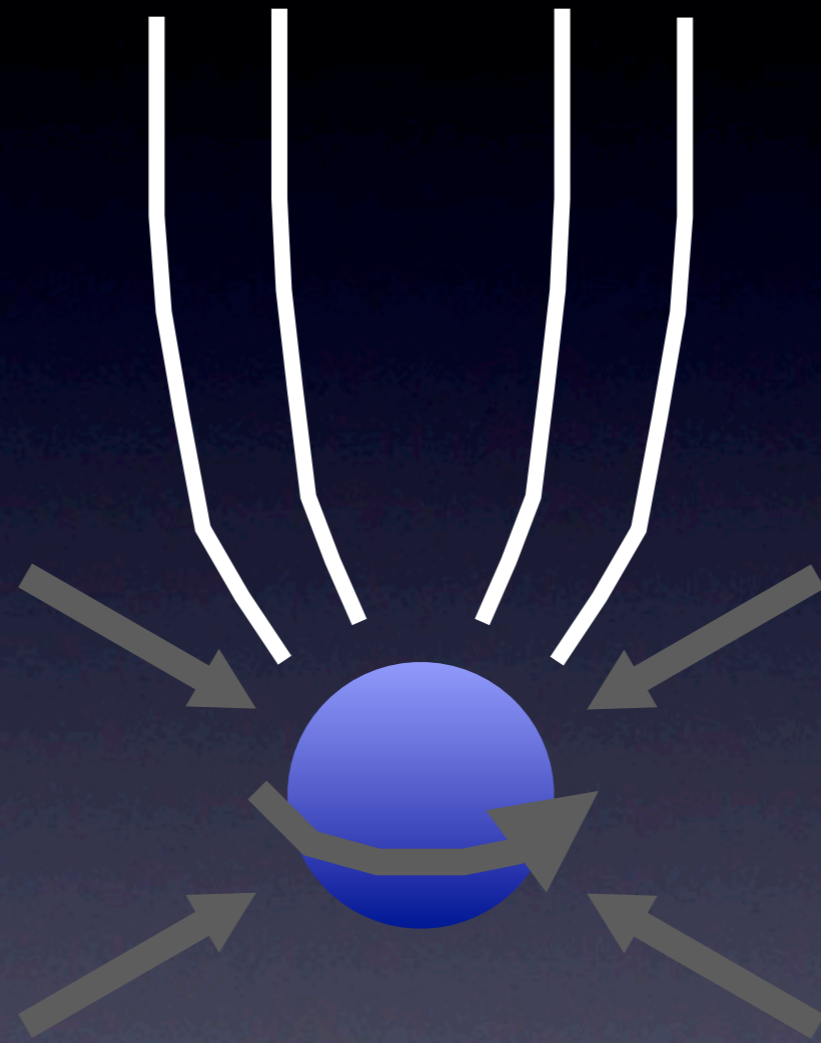
Credit: Nicolle Rager Fuller/NSF

Gamma-ray burst

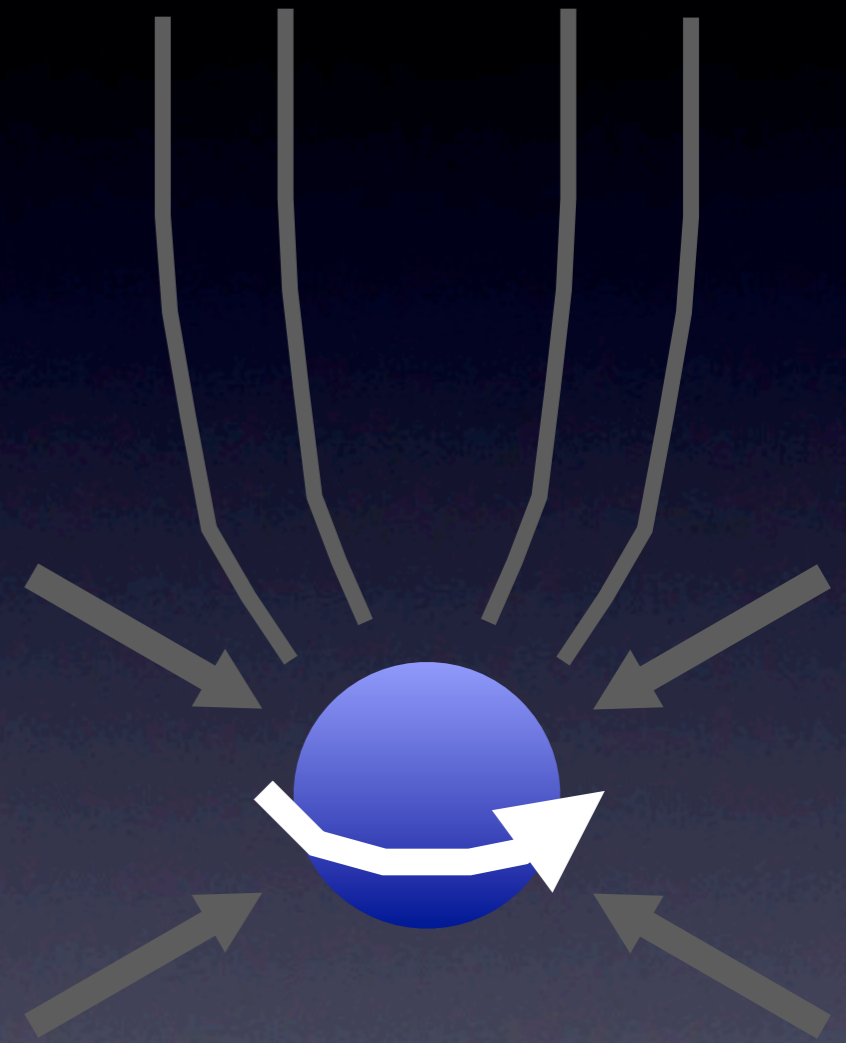
Ingredients



Accretion

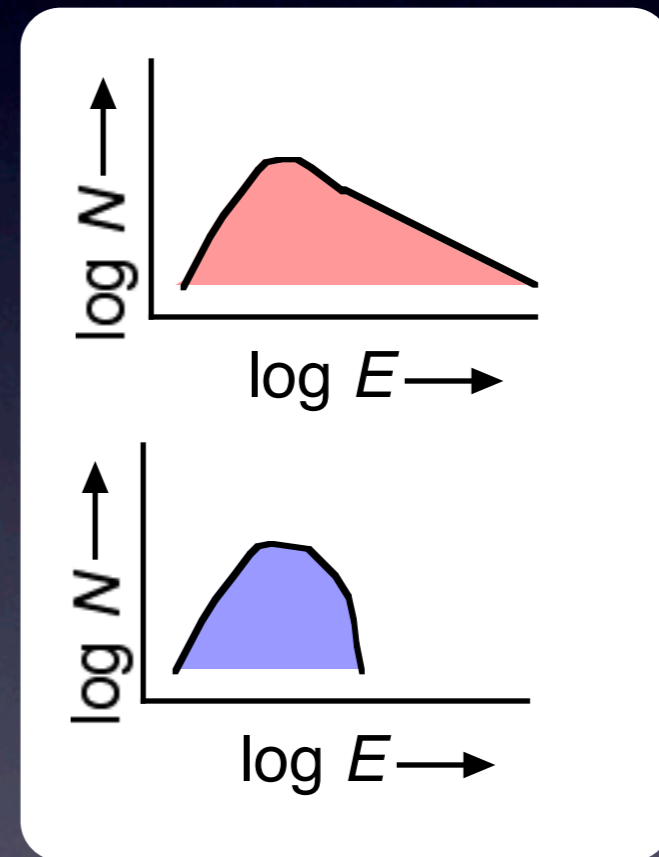
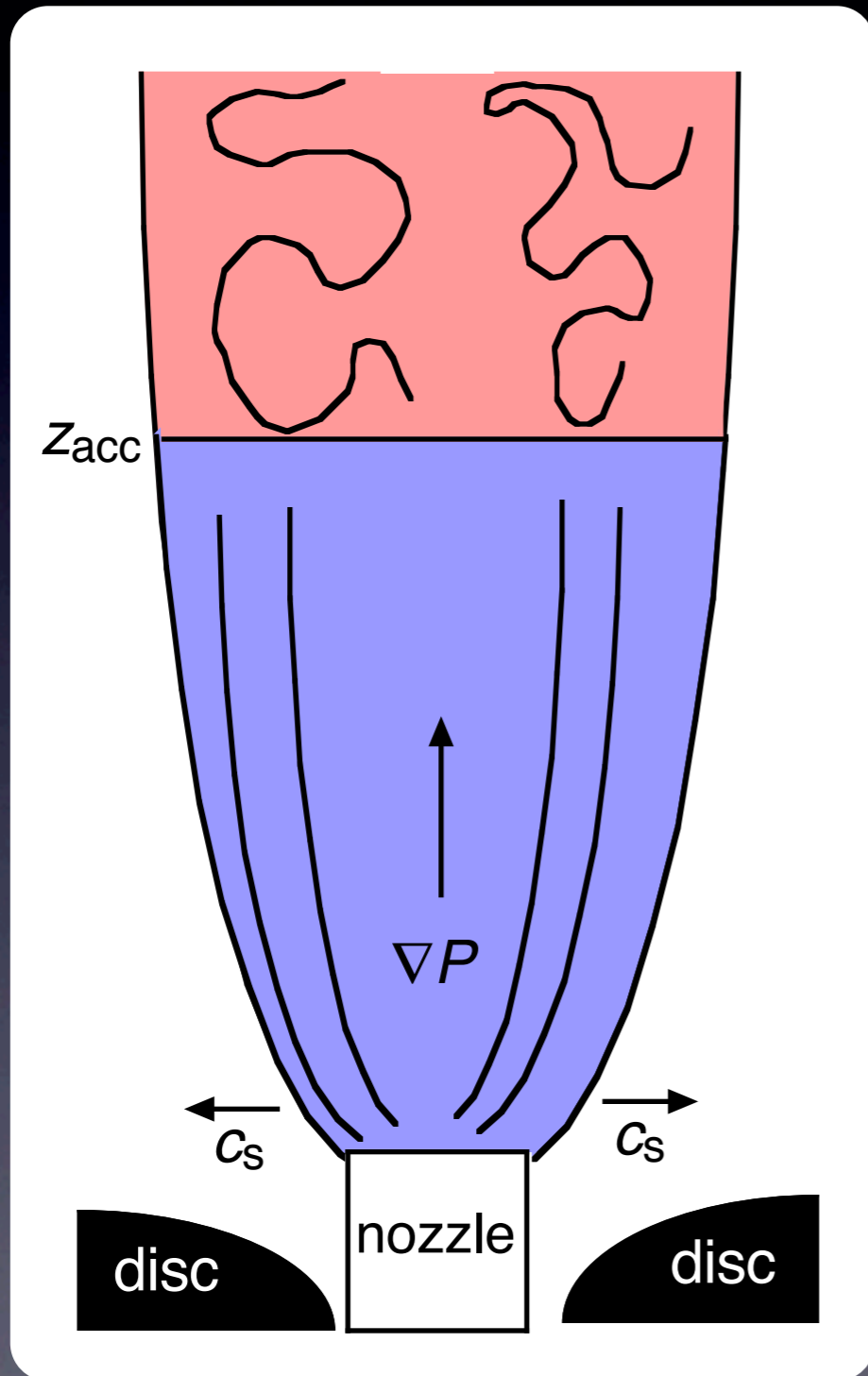


Magnetic fields

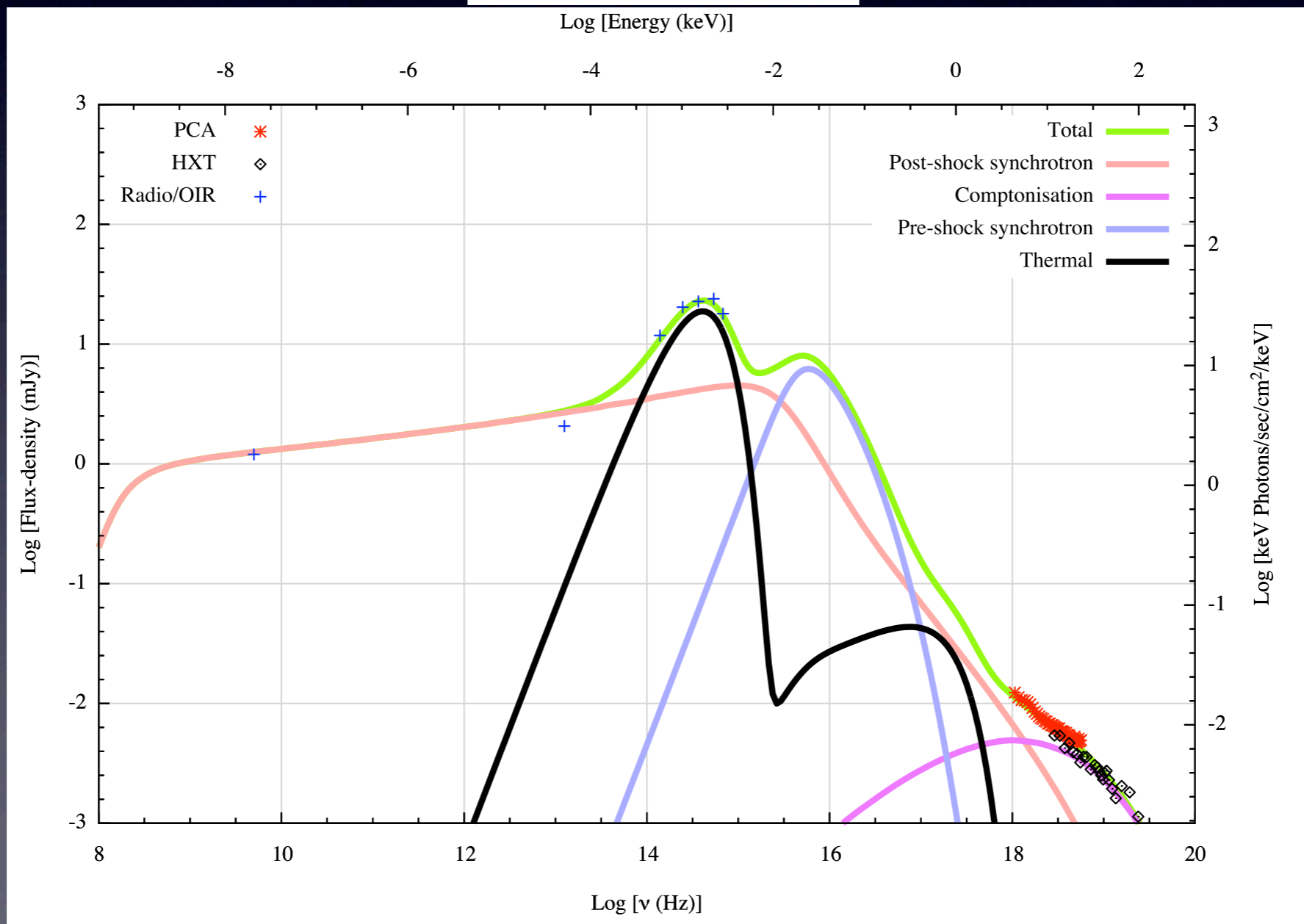
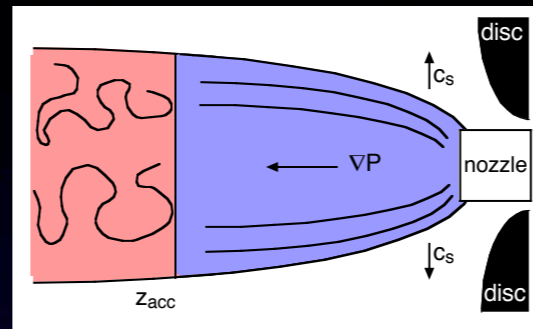
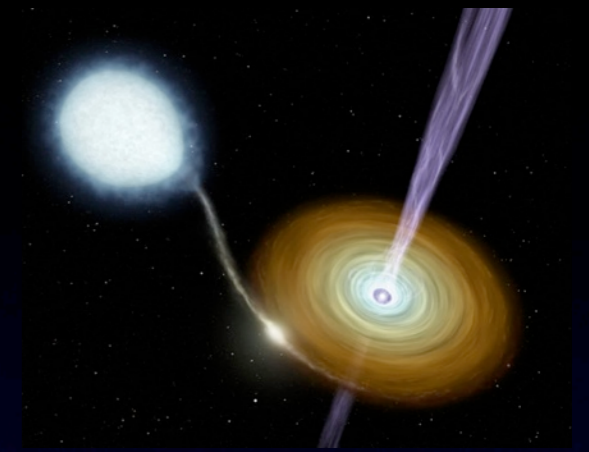


Rotation

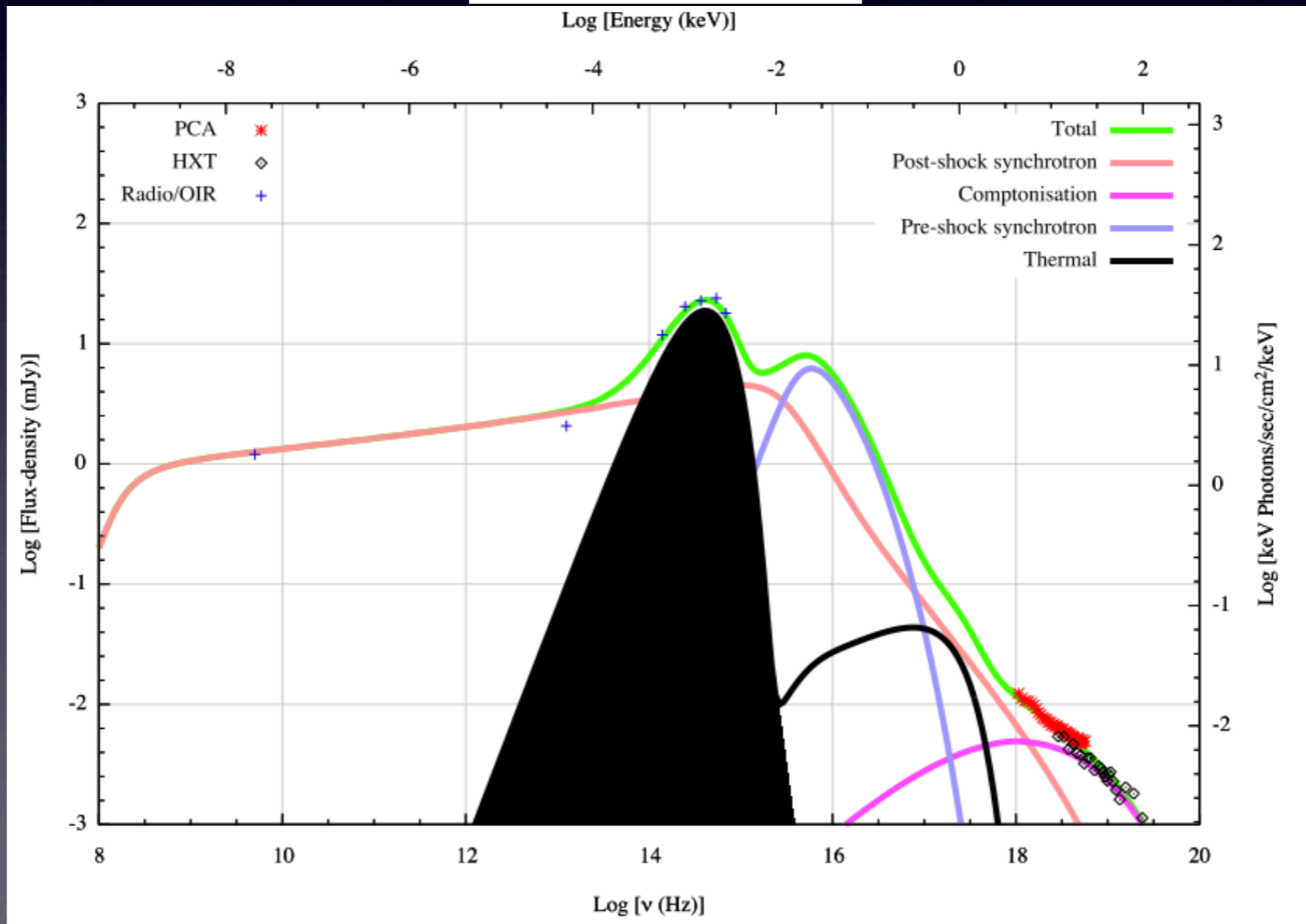
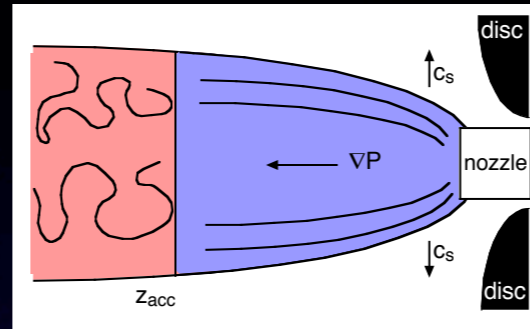
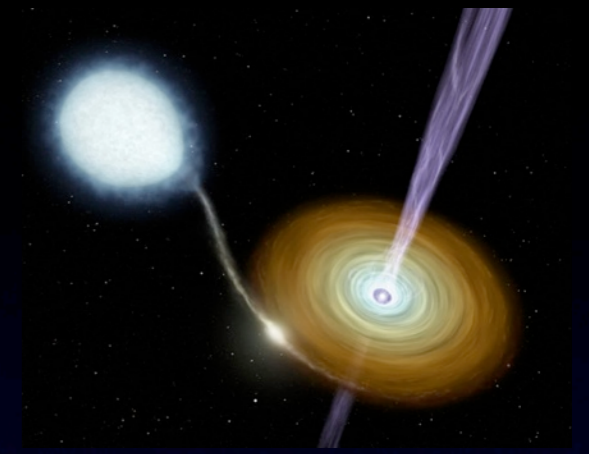
Acceleration region



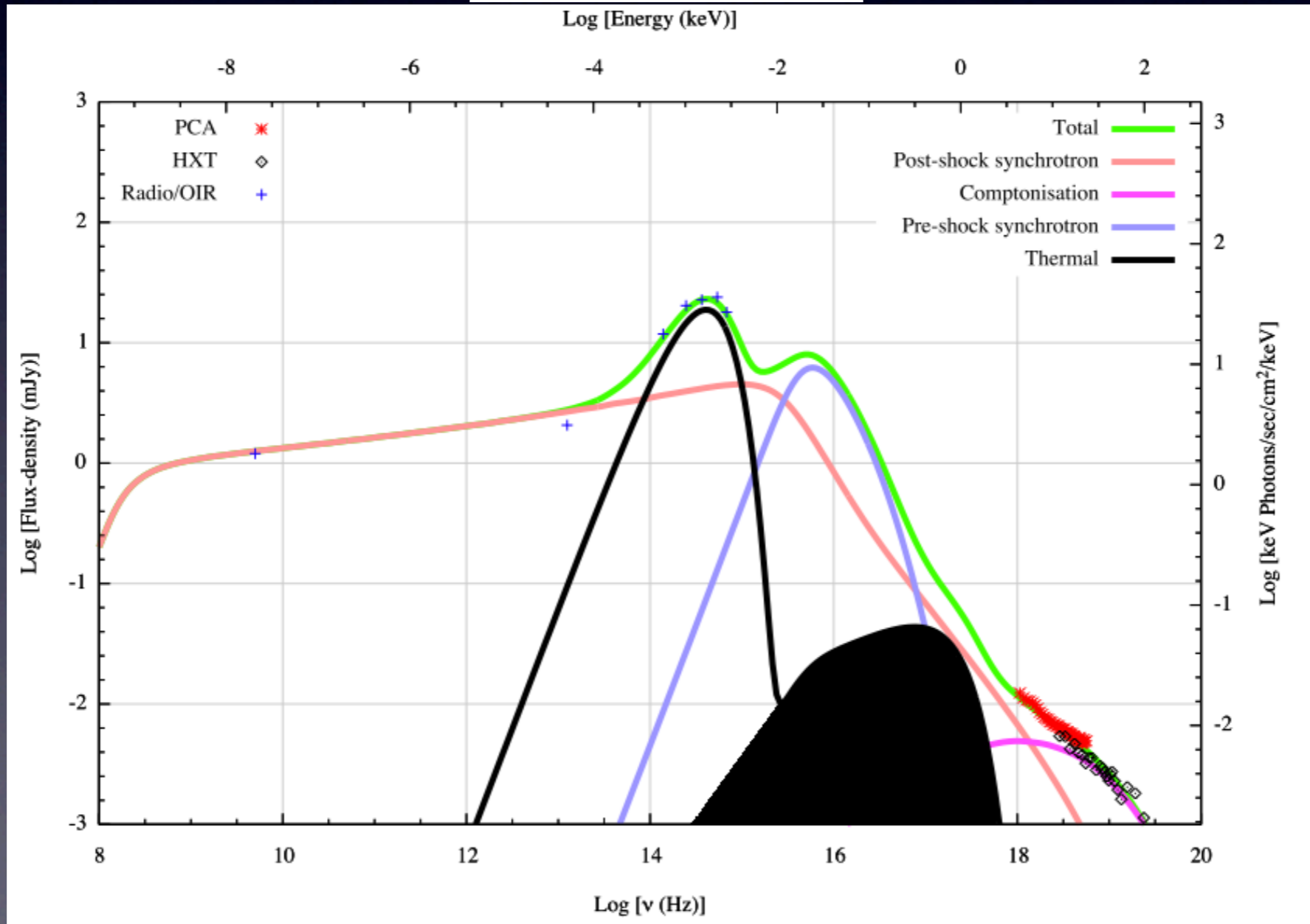
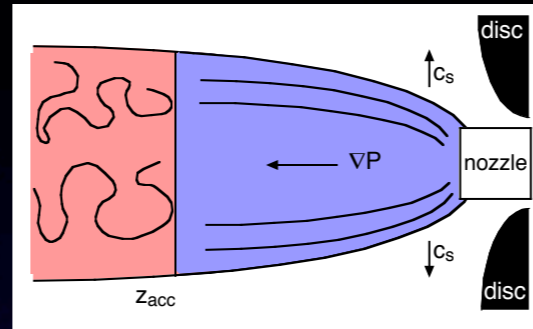
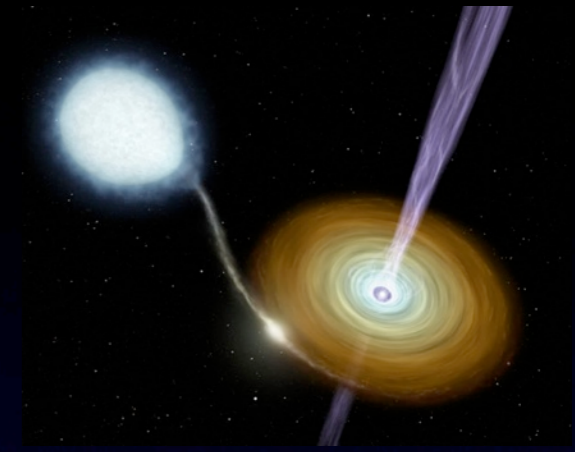
Spectral fitting



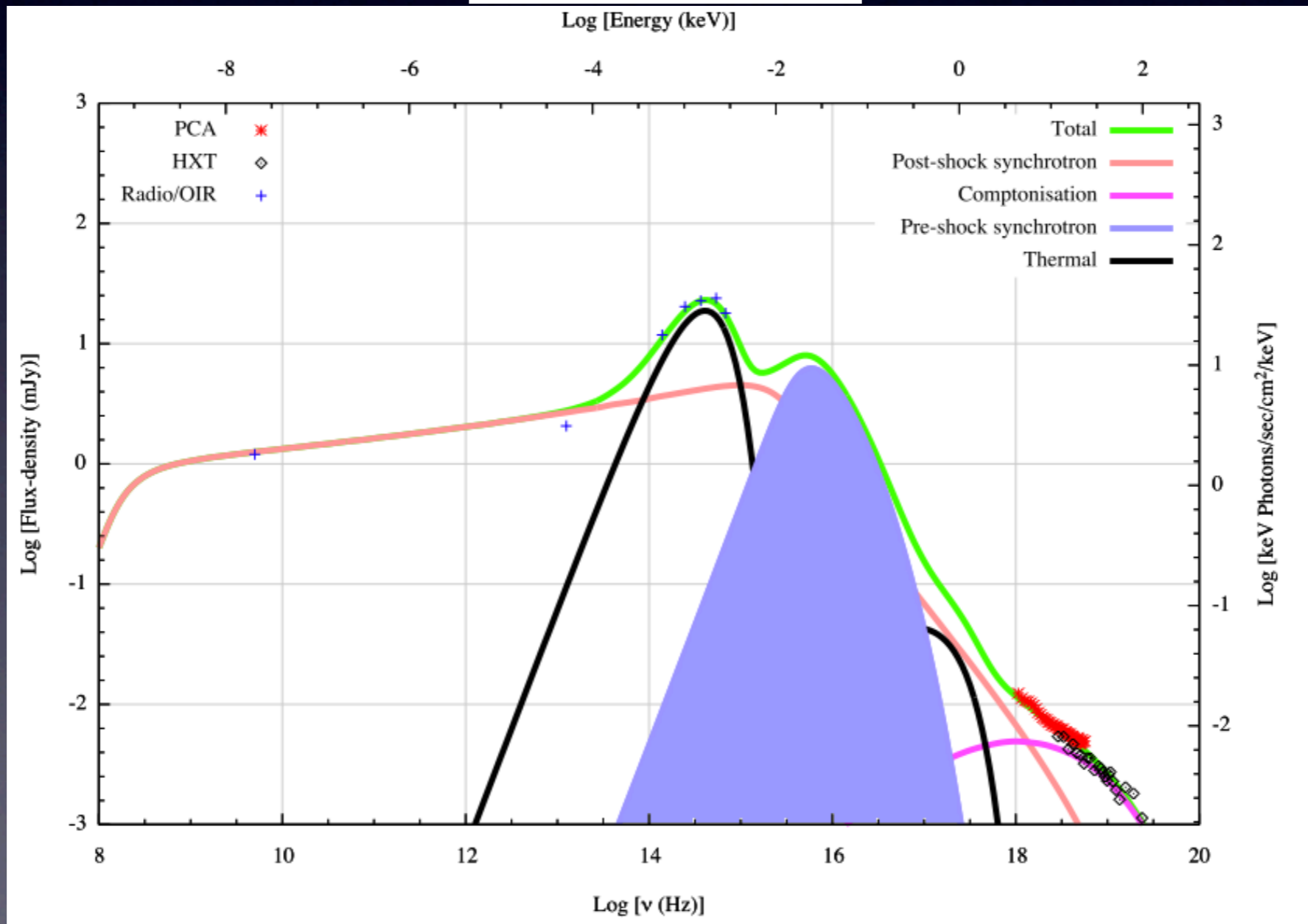
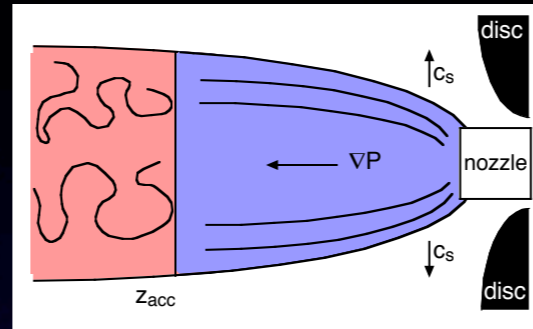
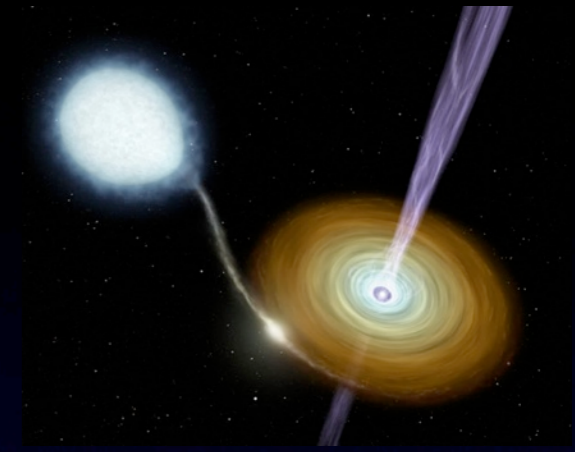
Spectral fitting



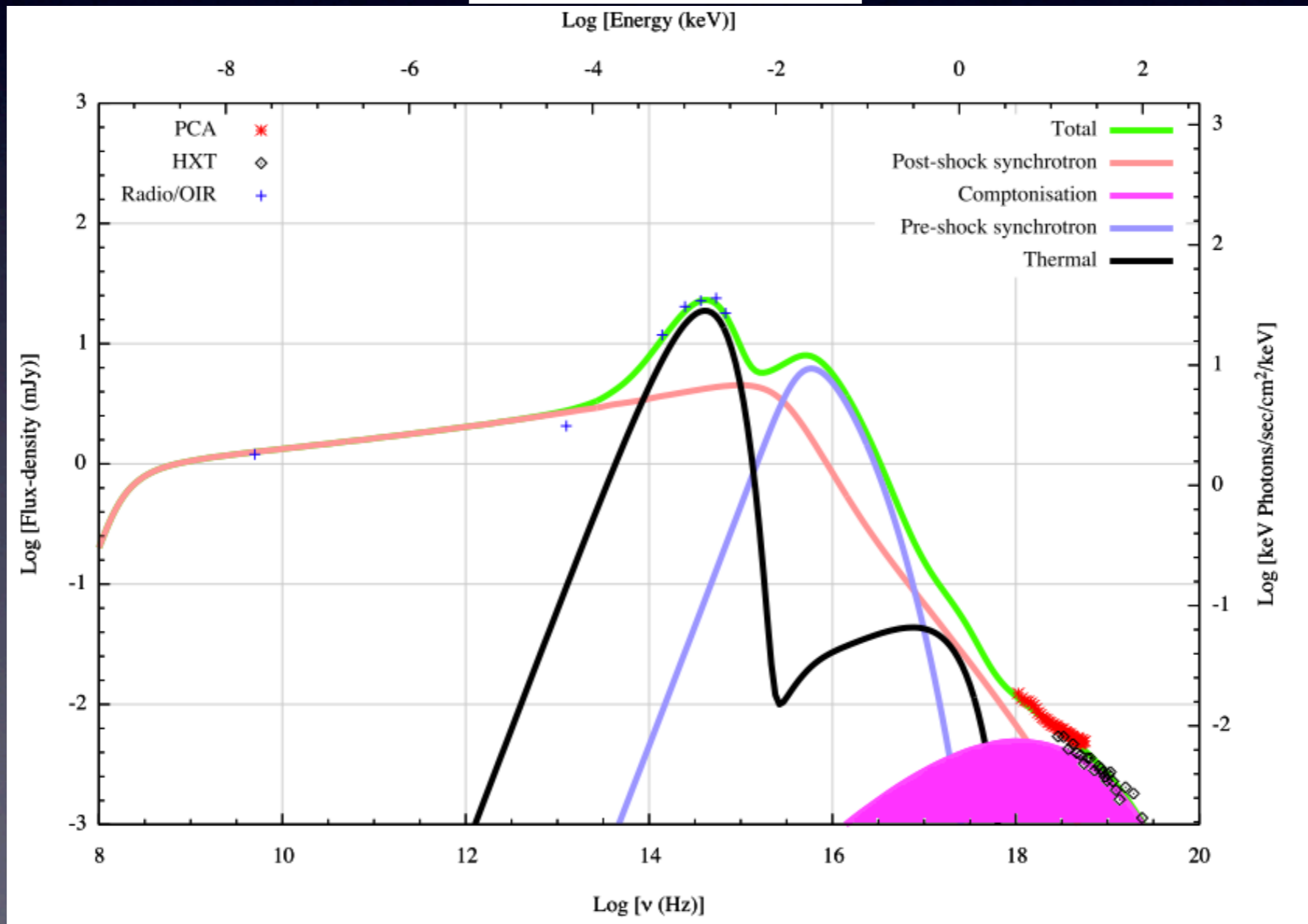
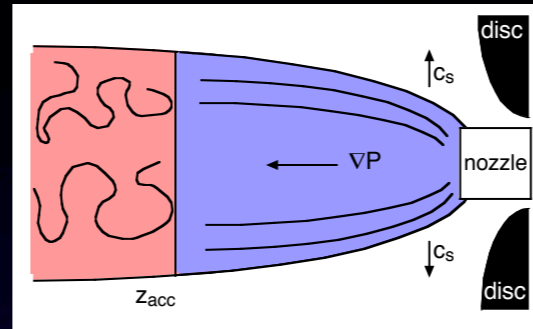
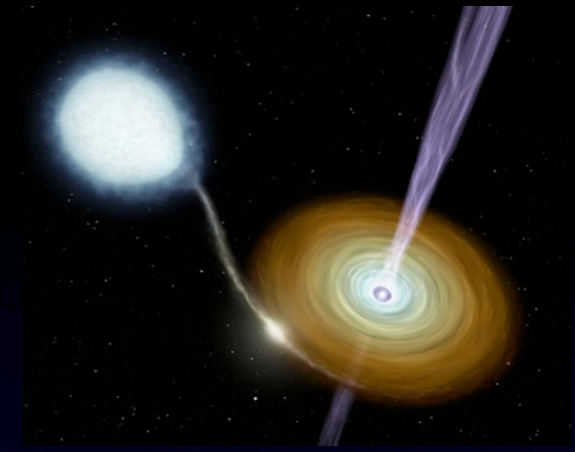
Spectral fitting



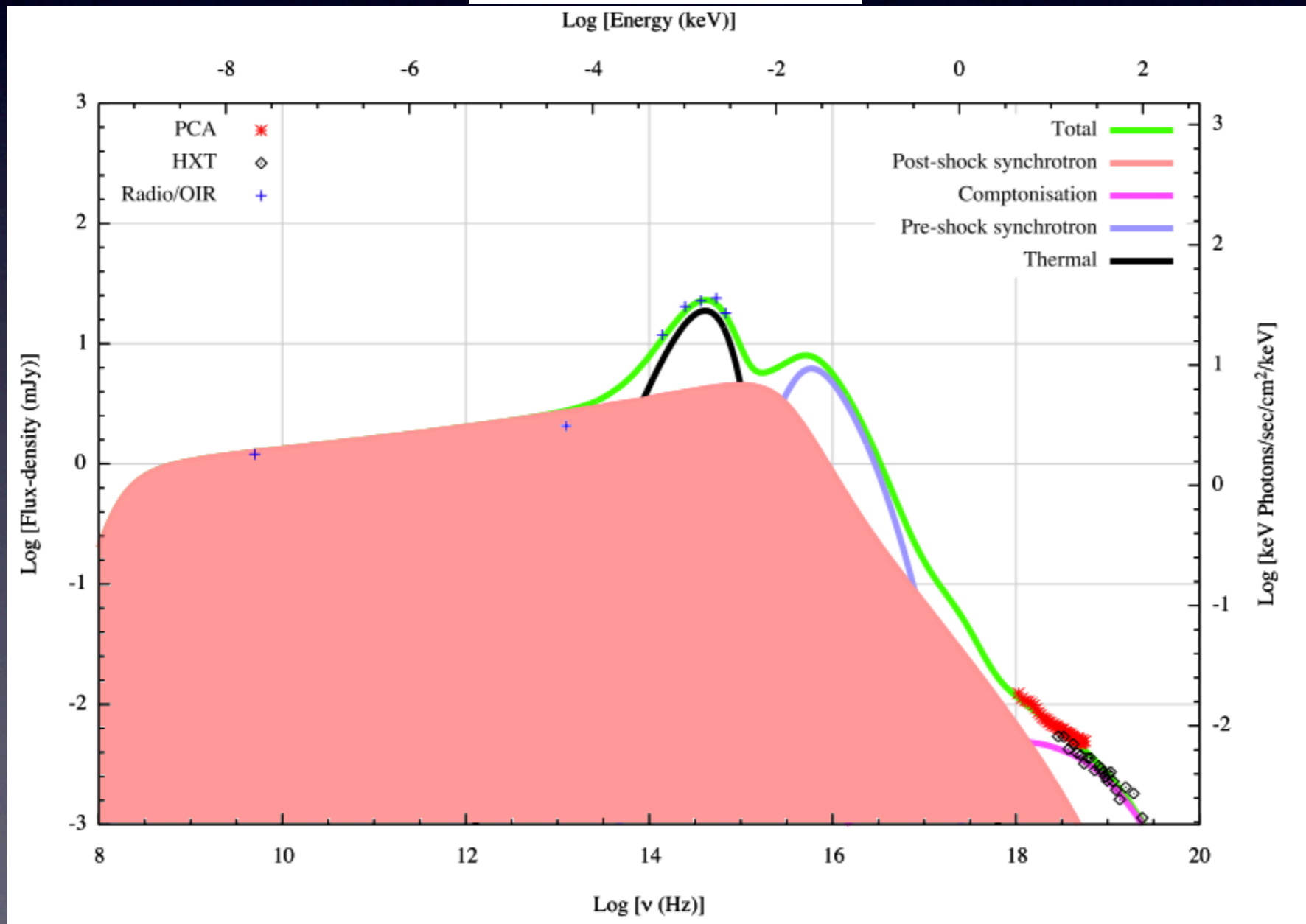
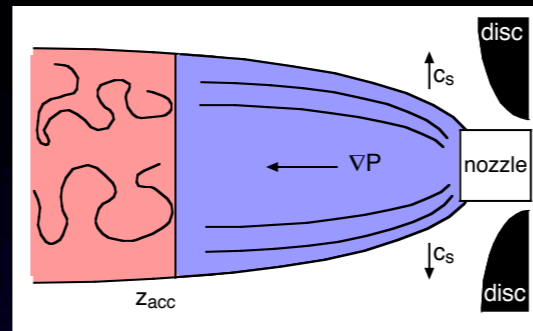
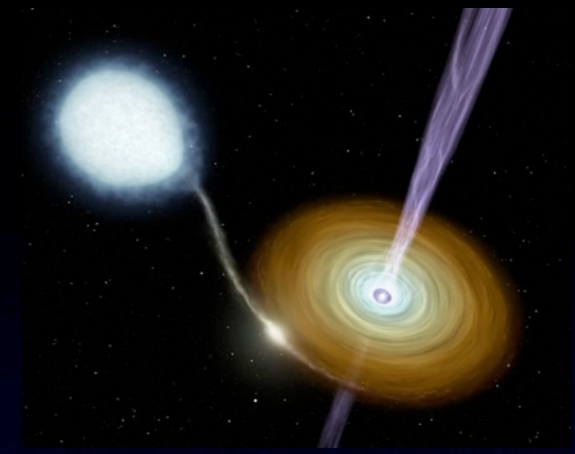
Spectral fitting



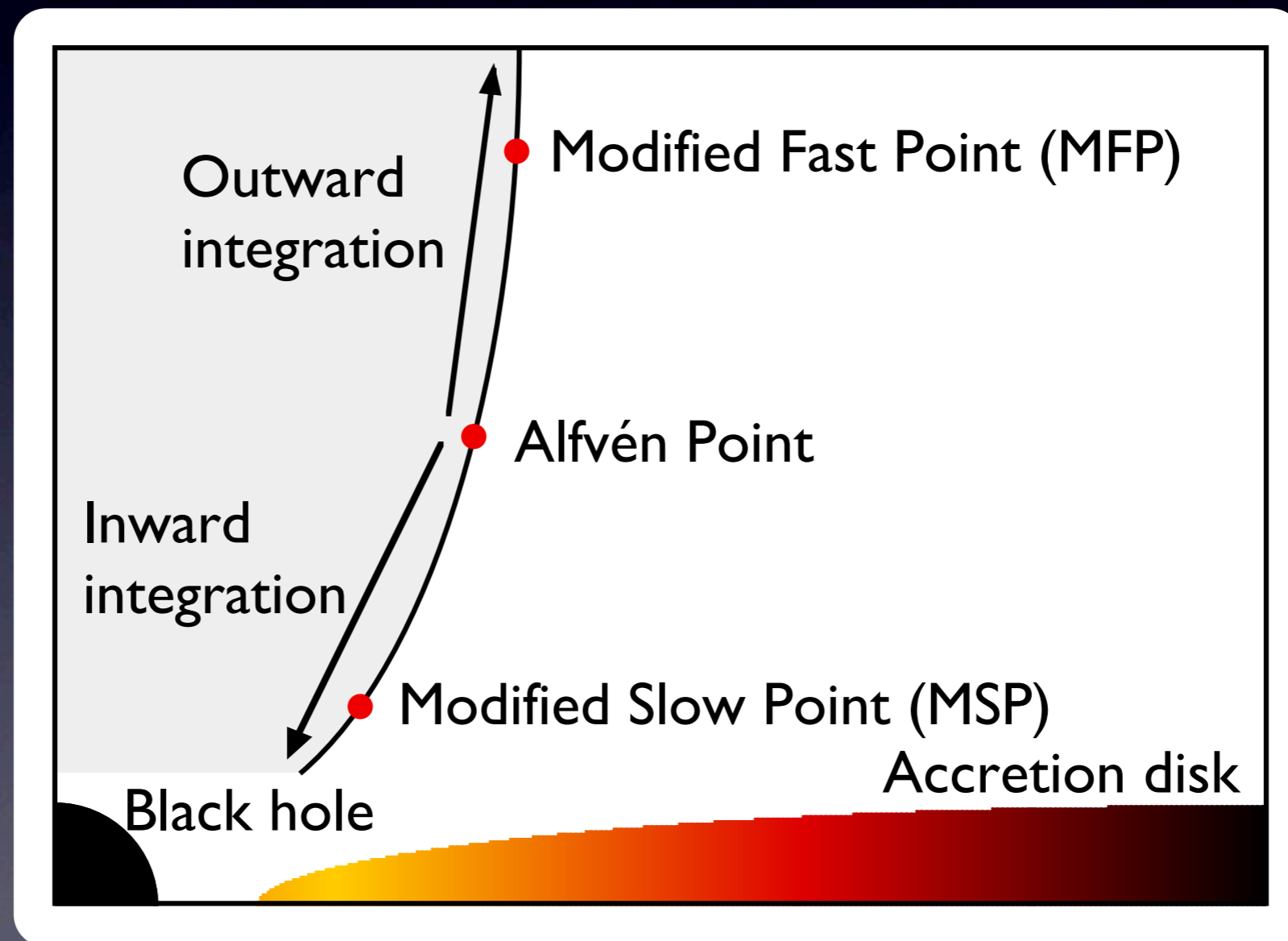
Spectral fitting



Spectral fitting

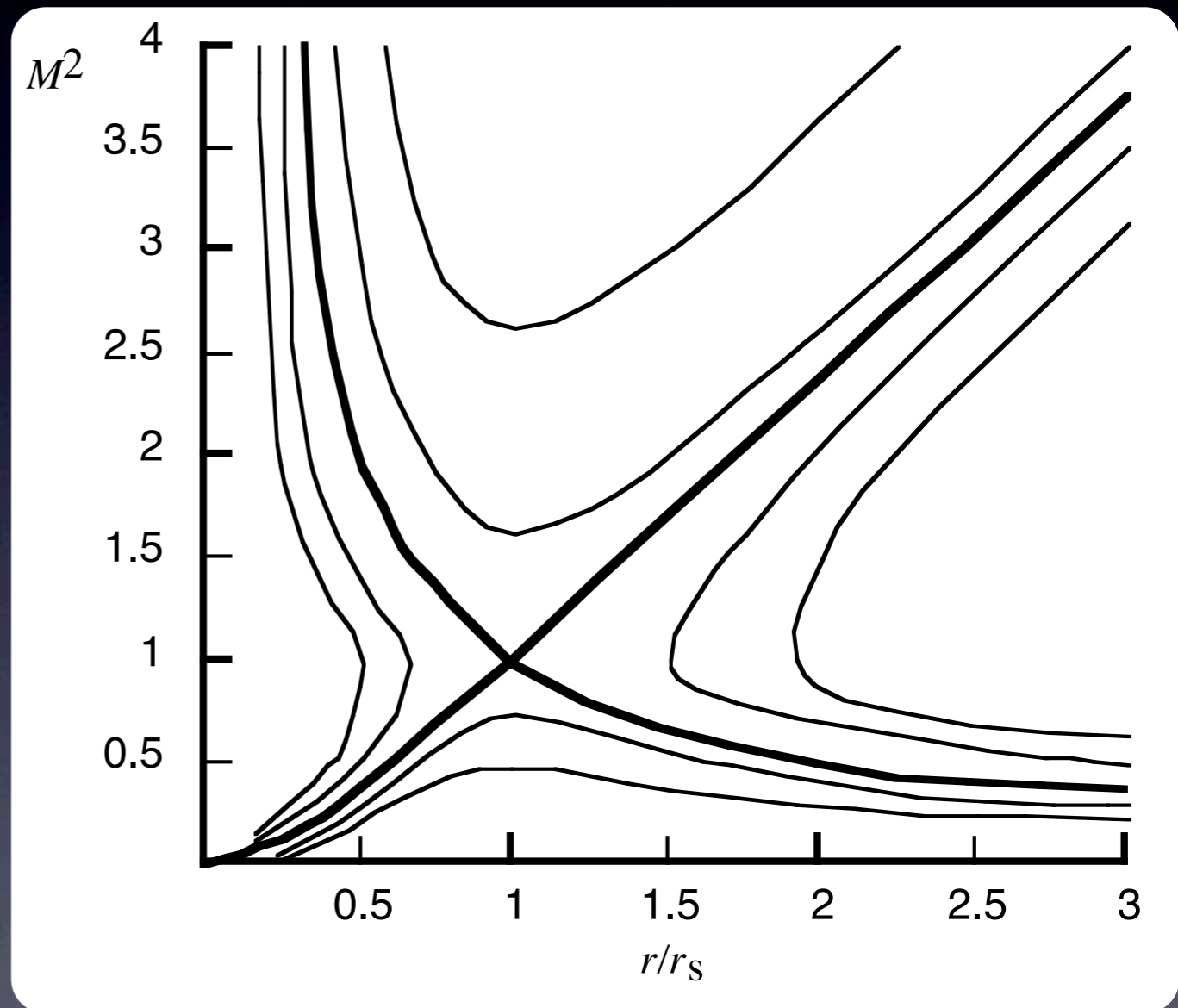


Critical surfaces

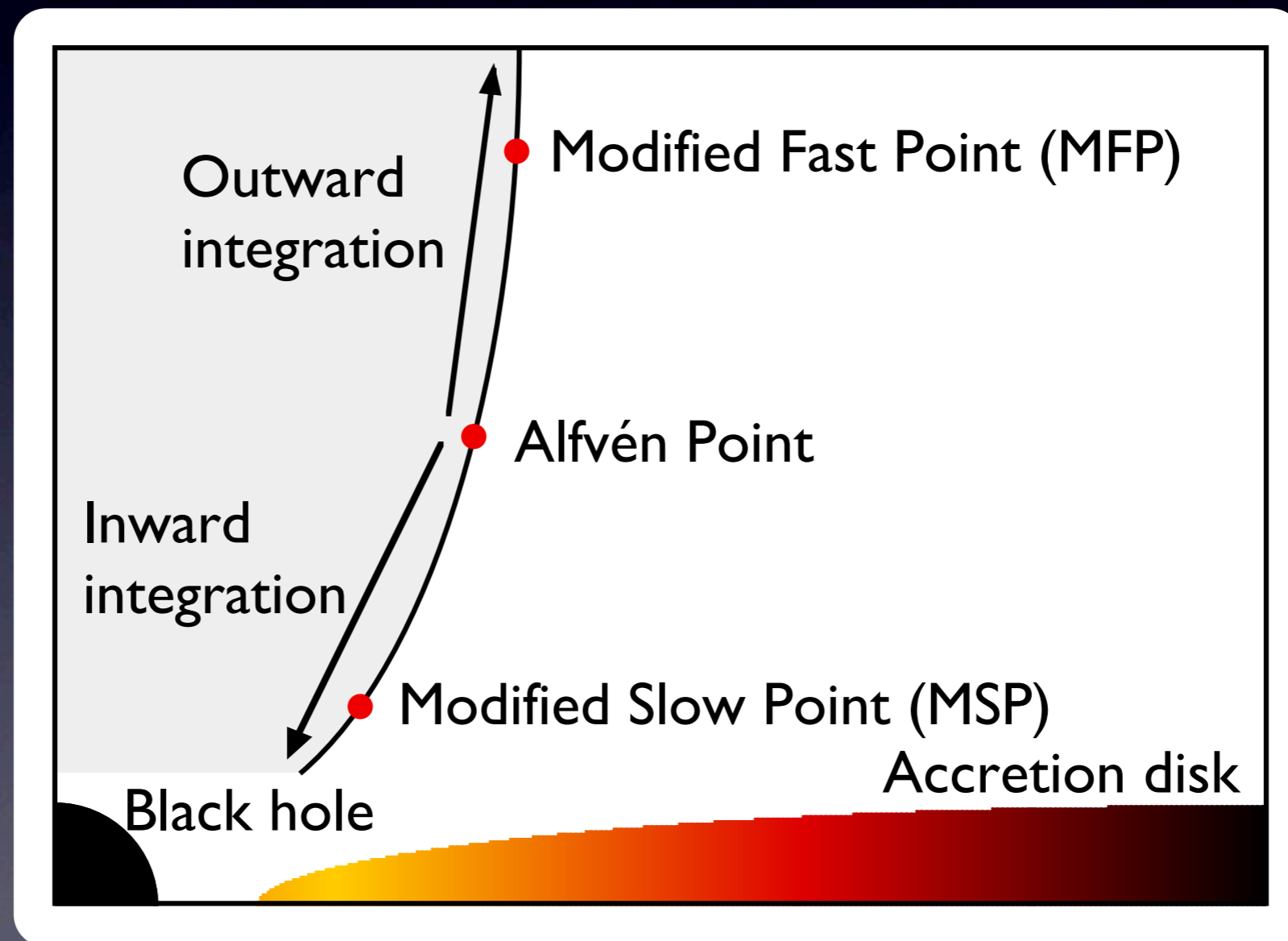


Parker wind

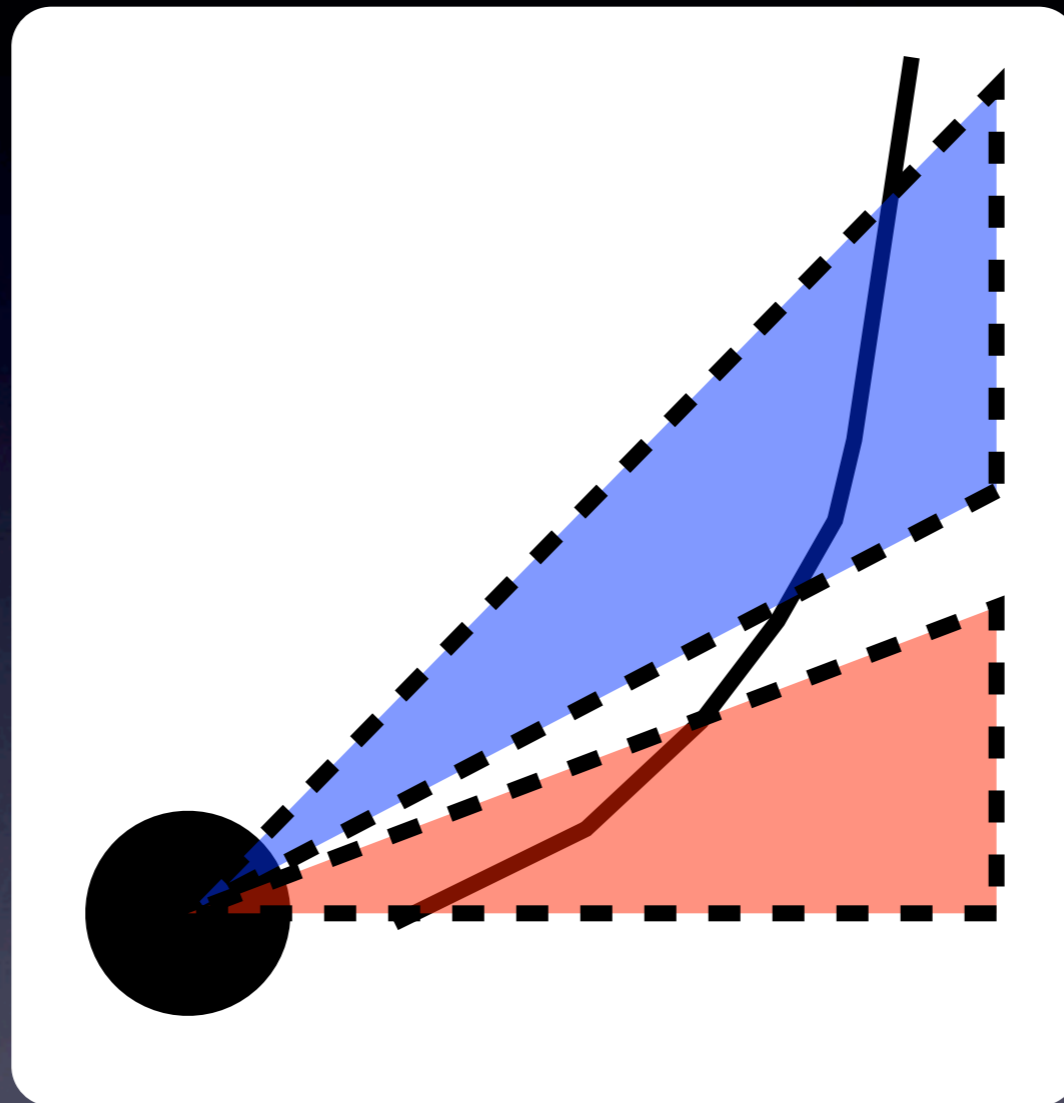
$$\frac{dv^2}{dr} = -\frac{2GM}{r^2} \left[\frac{1 - \left(\frac{2c_s^2 r}{GM} \right)}{1 - \frac{c_s^2}{v^2}} \right]$$



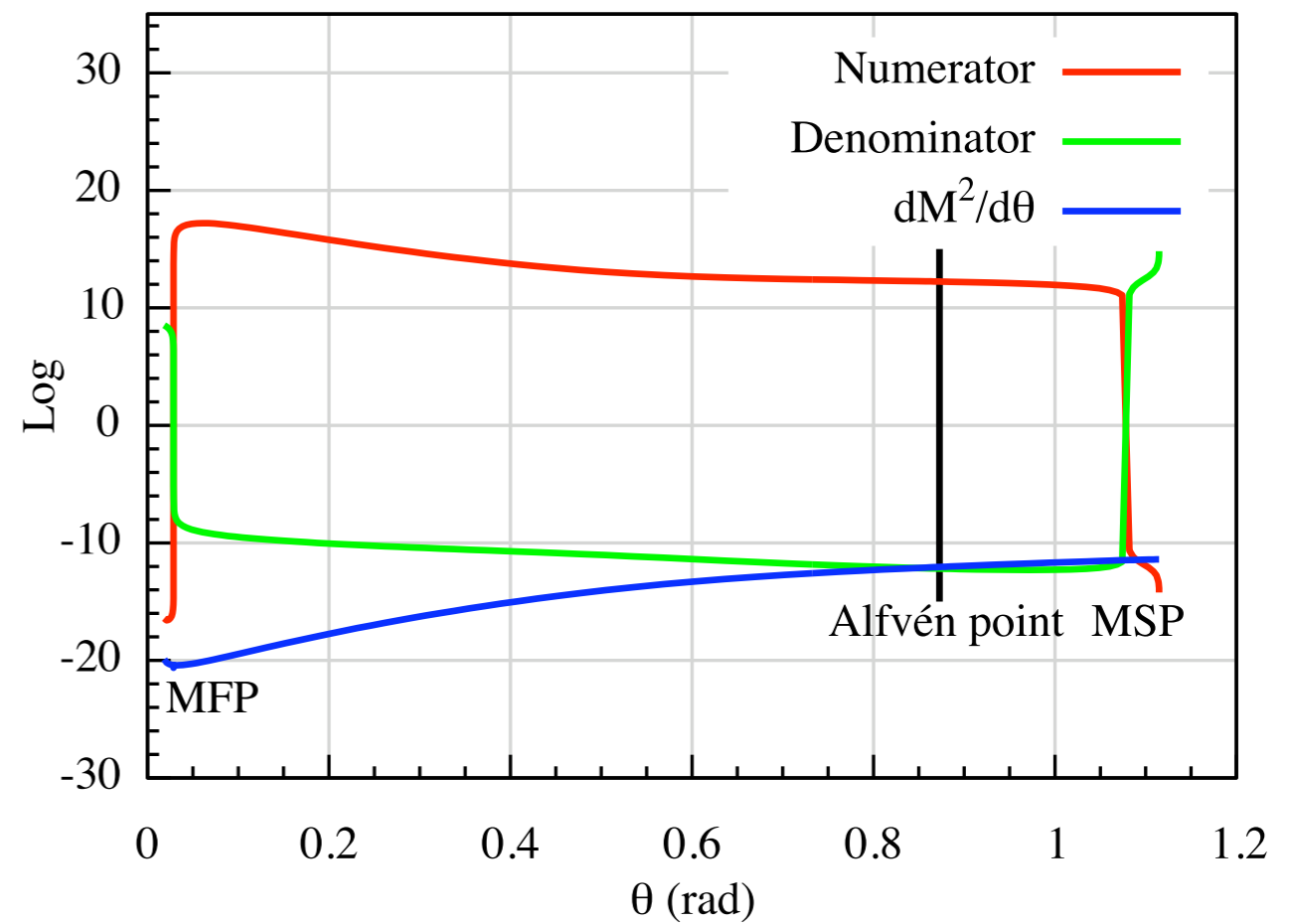
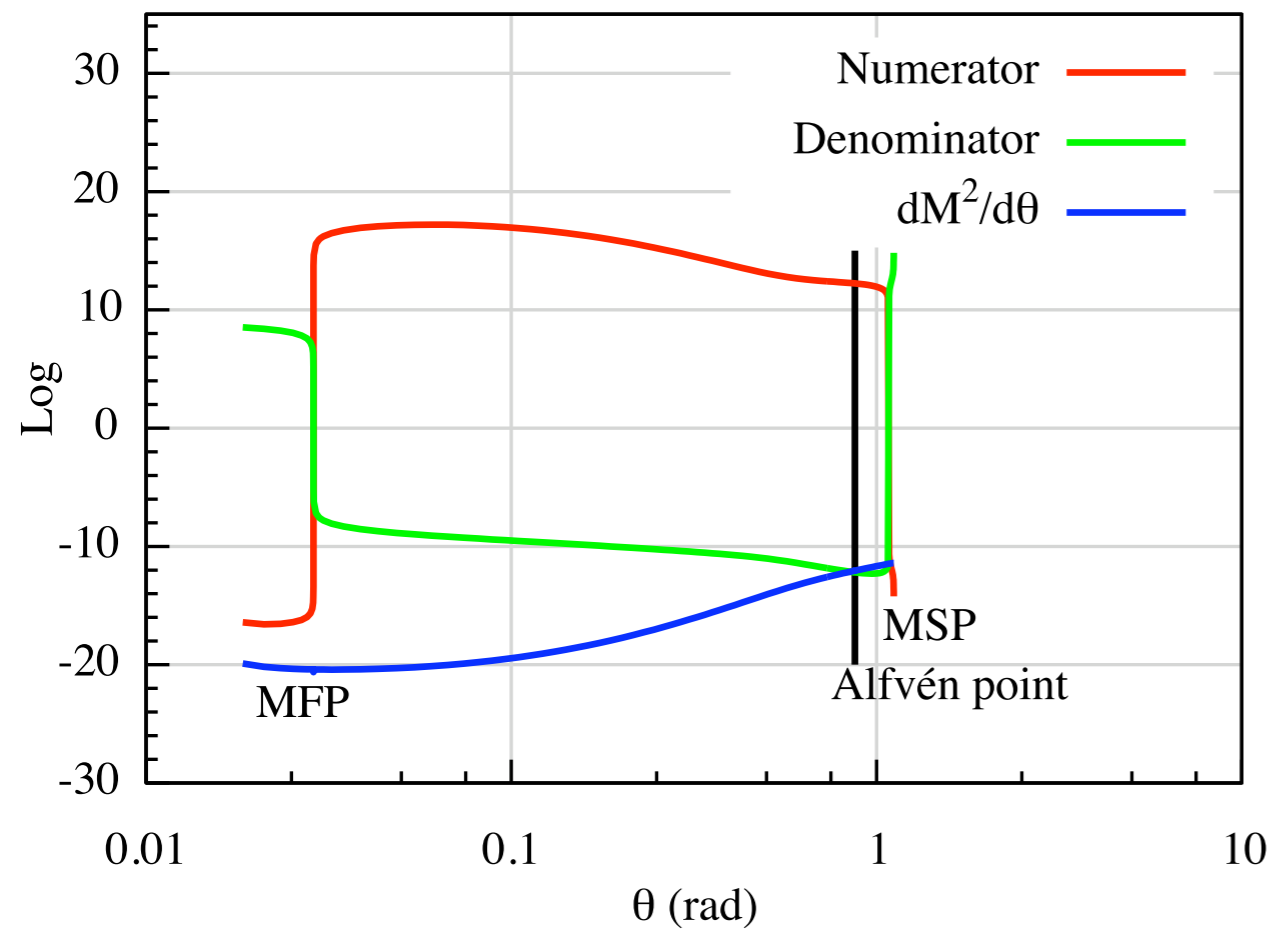
Singular surfaces



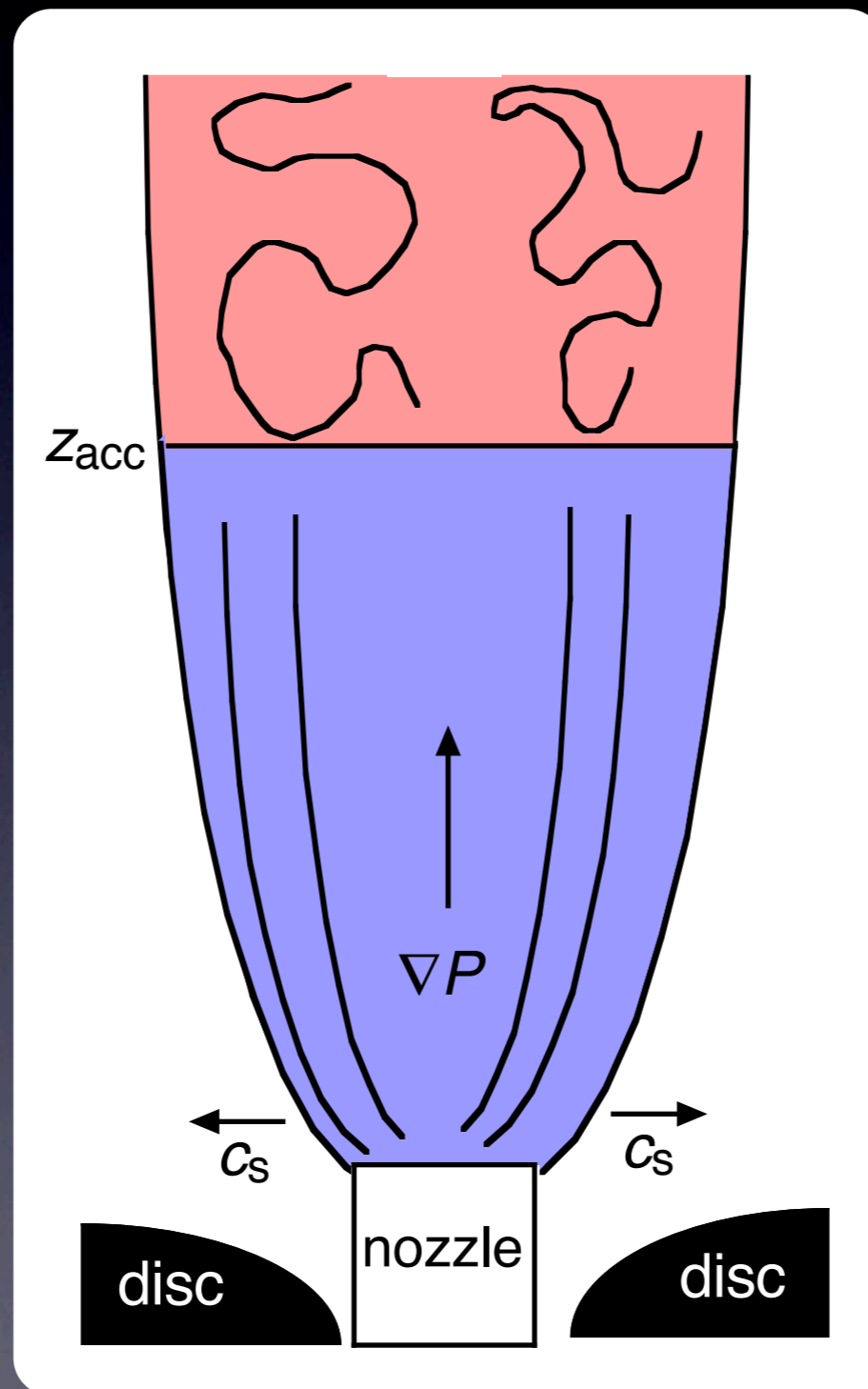
Combined solution



Critical solution



Connections



Conclusions

- The shock location helps us understand the spectrum
- Self-consistent model to determine shock location

